

SPT 0538–50: PHYSICAL CONDITIONS IN THE INTERSTELLAR MEDIUM OF A STRONGLY LENSED DUSTY STAR-FORMING GALAXY AT $z = 2.8$

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ABSTRACT

We present observations of SPT-S J053816–5030.8, a gravitationally lensed dusty star-forming galaxy (DSFG) at $z = 2.7817$ that was first discovered at millimeter wavelengths by the South Pole Telescope. SPT 0538–50 is typical of the brightest sources found by wide-field millimeter-wavelength surveys, being lensed by an intervening galaxy at moderate redshift (in this instance, at $z = 0.441$). We present a wide array of multi-wavelength spectroscopic and photometric data on SPT 0538–50, including data from ALMA, *Herschel* PACS and SPIRE, *Hubble*, *Spitzer*, the Very Large Telescope, ATCA, APEX, and the Submillimeter Array. We use high-resolution imaging from the *Hubble Space Telescope* to de-blend SPT 0538–50, separating DSFG emission from that of the foreground lens. Combined with a source model derived from ALMA imaging (which suggests a magnification factor of 21 ± 4), we derive the intrinsic properties of SPT 0538–50, including the stellar mass, far-IR luminosity, star formation rate, molecular gas mass, and—using molecular line fluxes—the excitation conditions within the interstellar medium. The derived physical properties argue that we are witnessing compact, merger-driven star formation in SPT 0538–50 similar to local starburst galaxies and unlike that seen in some other DSFGs at this epoch.

Key words: galaxies: evolution – galaxies: formation – galaxies: high-redshift – galaxies: individual – gravitational lensing: strong – submillimeter: galaxies

Online-only material: color figures

1. INTRODUCTION

The population of high-redshift, luminous dusty starbursts known as “submillimeter galaxies” was first discovered in 1998 with the SCUBA instrument on the James Clerk Maxwell telescope (Holland et al. 1999). Since then, a picture has

emerged of a population of massive ($M_* \sim 10^{11} M_\odot$; Hainline et al. 2011; Michałowski et al. 2012) and gas-rich ($M_{\text{gas}} > 10^{10} M_\odot$; Frayer et al. 1998; Neri et al. 2003; Greve et al. 2005; Bothwell et al. 2013) galaxies with IR luminosities that imply prodigious star formation rates (SFRs; $100 \text{ s to } > 1000 M_\odot \text{ yr}^{-1}$; Kovács et al. 2006). It is clear that they represent a population of

significant astrophysical importance with properties that remain a considerable challenge for conventional galaxy evolution models (Baugh et al. 2005; Lacey et al. 2010; Benson 2012; Hayward et al. 2013).

The advent of millimeter/submillimeter instruments capable of surveying large areas of sky, including *Herschel*-SPIRE (Pilbratt et al. 2010; Griffin et al. 2010) and the South Pole Telescope (SPT; Carlstrom et al. 2011), has opened new vistas for the study of these submillimeter bright systems (Vieira et al. 2010; Negrello et al. 2010). The brightest of these submillimeter sources, with $S_{850\mu\text{m}} > 100$ mJy, have such low sky density ($< 1 \text{ deg}^{-2}$) that large area surveys are the only way to detect them in statistically significant numbers. Follow-up imaging in the submillimeter at high spatial resolution ($< 1''$) has confirmed the initial predictions that these objects represent a population of lensed, dusty star-forming galaxies (DSFGs) magnified by gravitational lensing (Negrello et al. 2010; Vieira et al. 2013; Hezaveh et al. 2013).

In recent years, the study of lensed starburst systems at high redshift has offered invaluable insight into the behavior and physics of these dust-obscured systems. Strong gravitational lensing provides an effective boost in both sensitivity and angular resolution, allowing highly detailed analyses of lensed systems that would not be possible in their unlensed counterparts.

The $z = 2.3$ galaxy SMMJ2135–0102—lensed by a factor of ~ 30 —has been extensively studied with the lensing magnification allowing detailed observations of a large rotating molecular disk (Swinbank et al. 2011), a dusty IR spectral energy distribution (SED; Ivison et al. 2010), and a well-constrained CO spectral line energy distribution (SLED) implying interstellar medium (ISM) excitation conditions similar to local starbursts (Danielson et al. 2011). Fu et al. (2013) present a comprehensive analysis of an H-ATLAS-selected lensed DSFG at $z = 3.2$, a system for which the lensing magnification has revealed a highly complex, clumpy structure and an obscured UV-to-far IR SED. In addition, Wardlow et al. (2013) present 13 candidate gravitationally lensed DSFGs from the *Herschel* Multi-tiered Extragalactic Survey (HERMES), finding that their intrinsic submillimeter fluxes typically lie below the *Herschel* selection limit and, as such, represent systems that can only be studied via gravitational lensing.

SPT-S J053816–5030.8 (hereafter SPT 0538–50) is the first lensed DSFG selected by the SPT survey to be characterized in detail. SPT 0538–50 is taken from the catalog of sources from the first 87 deg^2 in the SPT survey (Vieira et al. 2010) with single-dish submillimeter follow-up presented in Greve et al. (2012). It was targeted for imaging with the Atacama Large Millimeter/submillimeter Array (ALMA)—these results, which were used to construct a physically motivated lens model of the system, were presented by Hezaveh et al. (2013). With this lens model in hand, we present a range of multi-wavelength observations of the source and derive the intrinsic luminosity, SFR, stellar mass, molecular gas mass, and physical conditions of this interesting galaxy.

The paper is organized as follows. In Section 2, we present our multiwavelength observations of SPT 0538–50, which we divide into imaging (Section 2.1) and spectroscopy (Section 2.2). In Section 3, we present our analysis and in Section 4, we discuss the implications for the physical properties of SPT 0538–50. We conclude in Section 5. All magnitudes in this work are given in the Vega system. Throughout this work we adopt a cosmological model with $(\Omega_m, \Omega_\Lambda, H_0) = (0.27, 0.73, 71 \text{ km s}^{-1} \text{ Mpc}^{-1})$

and a Kroupa (2001) initial mass function (IMF; which is very similar to the Chabrier 2003 IMF; masses and luminosities derived using a Kroupa IMF can be converted to Chabrier by multiplying by a factor of 1.03).

2. OBSERVATIONS AND RESULTS

Since its discovery, SPT 0538–50 (ALMA co-ordinates R.A.: 05:38:16.83, decl.: 50:30:52.0) has been observed in a multitude of wavebands, from visible to radio wavelengths. The resulting datasets enable us to investigate the intrinsic physical properties of the ISM, while also constraining the stellar mass. Here we present the observational data collected for SPT 0538–50 starting with submillimeter and millimeter imaging data, moving on to NIR and optical imaging, and ending with millimeter-wavelength spectroscopy.

2.1. Imaging

Here we present observational details for the imaging data used in this work. Data from the SPT (Section 2.1.1), ALMA (Section 2.1.2), Australia Telescope Compact Array (ATCA; Section 2.1.3) and Atacama Pathfinder Experiment (APEX; Section 2.1.4) have been previously published by our group (see the relevant sections for references). We also present unpublished data from *Herschel* (Section 2.1.5), *Spitzer*/IRAC (Section 2.1.6), VLT/ISAAC (Section 2.1.7) and *HST*/WFC3 (Section 2.1.8).

2.1.1. SPT

SPT 0538–50 was initially detected at 1.4 mm and 2.0 mm wavelengths by the SPT. Details of these observations are given by Vieira et al. (2010). An updated description of the SPT survey and point source catalog can be found in Mocanu et al. (2013).

2.1.2. ALMA

The ALMA observations are described in Hezaveh et al. (2013)—here we provide a summary. SPT 0538–50 was imaged by ALMA in Cycle 0 (2011.0.00958.S; PI: D. Marrone) on 2011 November 16 for a total on-source integration time of 60 s. Sixteen antennas were incorporated into the array at the time of observation. Baselines extended from 20–150 k λ , creating a beam with a full width at half maximum (FWHM) of $\sim 1''.4$. The Band 7 local oscillator was tuned to 343.8 GHz with four correlator spectral windows configured to receive 1.875 GHz wide bands in two polarizations each at central frequencies of 336.8, 338.675, 348.925, and 350.8 GHz, for a total effective bandwidth of 7.5 GHz. Bandpass calibration was performed using observations of quasar J0538-440, while complex gain solutions were calculated using quasar J0519-4546. Absolute flux calibration was determined using observations of Callisto.

To register the ALMA data with images at other wavelengths, the primary phase calibrator was chosen to be registered to the International Celestial Reference Frame (ICRF; Ma et al. 1998). The calibrator and an additional test source, J0540-5418, are both registered to the ICRF with the precision of a few milliarcseconds (Petrov et al. 2011). This test source was observed less than three minutes after SPT 0538–50 to verify the astrometric calibration. Transferring phase and amplitude calibration from the phase calibrator J0519-4546, 9''.1 away, recovers the position of the test source to within 0''.09, suggesting a similar precision for SPT 0538–50, which is 5''.6 from the calibrator and 3''.8 from the test source.

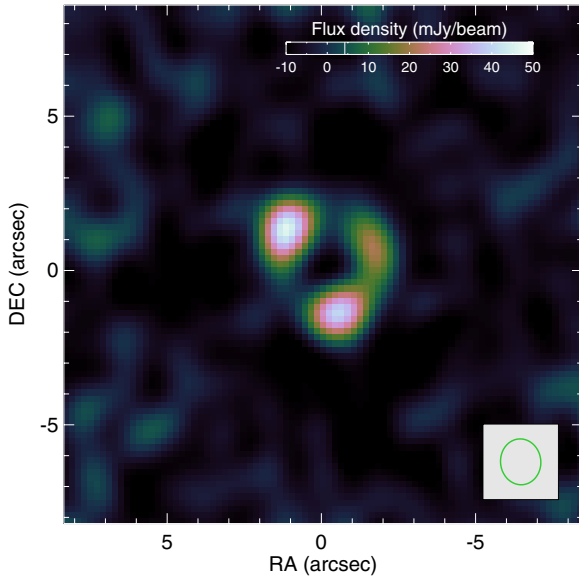


Figure 1. ALMA Band 7 ($=349.8$ GHz) continuum image of SPT 0538–50 taken with an integration time of 60 s. The synthesized beam FWHM ($1''.3 \times 1''.5$, at P.A. $= 5^\circ$) is shown in the lower right inset. The observations resolve the submillimeter emission into a clear Einstein ring with a radius of $1''.8$, conclusively demonstrating the lensed origin of SPT 0538–50.

(A color version of this figure is available in the online journal.)

Our ALMA Band 7 compact-array observations (see Figure 1) have an rms noise of 1.0 mJy beam $^{-1}$. SPT 0538–50 is detected with a signal-to-noise ratio (S/N) of ~ 50 , and is resolved into a clear Einstein ring, which conclusively indicates a lensed origin for the source.

2.1.3. ATCA

SPT 0538–50 was observed with the ATCA on 2012 July 31 and was well detected in both $^{12}\text{CO}(1\rightarrow 0)$ and the 35 GHz continuum. The final data resulted in a continuum level of 130 ± 20 μJy , a continuum rms of 20 μJy at 35 GHz, and an rms of 0.60 mJy beam $^{-1}$ per 2 MHz channel at 30.6 GHz (near the location of the CO emission line). Further details are reported in Aravena et al. (2013).

2.1.4. APEX LABOCA and SABOCA

SPT 0538–50 was imaged with the APEX Submillimetre APEX Bolometer Camera (SABOCA) and the Large APEX Bolometer Camera (LABOCA) instruments in May of 2010 in excellent weather conditions (see Siringo et al. 2009 for instrument details). On-source integration times for SABOCA 350 μm and LABOCA 870 μm observations were approximately 3 hr and 1 hr, respectively, reaching respective noise levels of 30 mJy beam $^{-1}$, and 6 mJy beam $^{-1}$. SPT 0538–50 was detected with high significance in both bands, with resulting fluxes $S_{350\mu\text{m}} = 336 \pm 88$ mJy and $S_{870\mu\text{m}} = 125 \pm 7$ mJy (see Greve et al. 2012 for further details on these observations). SPT 0538–50 is unresolved in both bands, which have a beam FWHMs of $\sim 20''$ (at 870 μm) and $\sim 8''$ (at 350 μm).

2.1.5. Herschel PACS and SPIRE

The acquisition and reduction of data from both PACS and SPIRE (Griffin et al. 2010) are described extensively in Clements et al. (2010), Ibar et al. (2010), and Rigby et al. (2011).

The PACS data were acquired in programs OT1_jvieira_4 and OT1_dmarrone_1 using orthogonal scans centered on the

target, scanning at medium speed (i.e., with the telescope tracking at $20''$ s $^{-1}$) in “scan map” mode. Data were recorded simultaneously at 100 and 160 μm , and a total of 180 s of data were acquired in each program. Each scan comprises ten separate $3'$ strips, each offset orthogonally by $4''$. The data were then handled by a variant of the reduction pipeline presented in Ibar et al. (2010). The resulting noise levels were $\sigma \approx 4$ and 7 mJy at 100 and 160 μm , respectively (calculated using random aperture photometry).

The SPIRE data from program OT1_jvieira_4 consist of a triple repetition map with coverage complete to a radius of five arcmin from the nominal SPT position. The maps were produced via the standard reduction pipeline HIPE v9.0, the SPIRE Photometer Interactive Analysis package v1.7, and the calibration product v8.1. Photometry was extracted by fitting a Gaussian profile to the position of the SPT detection and the noise was estimated by taking the rms in the central five arcmin of the map, which was then added in quadrature to the absolute calibration uncertainty.

2.1.6. Spitzer IRAC

SPT 0538–50 was also observed at 3.6 μm and 4.5 μm by the Infrared Array Camera (IRAC; Fazio et al. 2004) on board the *Spitzer Space Telescope* (Werner et al. 2004) on 2009 August 2 as part of a large program to obtain follow-up imaging of the SPT sources (PID 60194; PI: Vieira). The basic calibrated data, pre-processed by the Spitzer Science Center’s standard pipeline, were combined into a re-sampled mosaic image using IRACproc (Schuster et al. 2006) as a wrapper to optimize the performance of the MOPEX software package (Makovoz & Marleau 2005). The resultant IRAC mosaic images have an angular resolution of $\sim 2''$, and a re-sampled pixel scale of $0''.6$ pixel $^{-1}$.

The foreground lens galaxy is well detected in both IRAC bands, with flux densities at 3.6 μm and 4.5 μm of $S_{3.6} = 0.38 \pm 0.01$ mJy and $S_{4.5} = 0.40 \pm 0.01$ mJy. Due to the small angular separation between the lensed image of SPT 0538–50 and the foreground lens, the (relatively) large IRAC beam causes emission from the two sources to be blended. In Section 3, we describe our de-blending method, which allows us to separate emission from SPT 0538–50 from that of the foreground lens.

2.1.7. VLT ISAAC

A deep *Ks*-band image of SPT 0538–50 was obtained using ISAAC (Moorwood et al. 1998), which was mounted on the UT1 (Antu) of the Very Large Telescope (VLT) at the European Southern Observatory (ESO) in Paranal, Chile. ISAAC is equipped with a 1024×1024 pixel Hawaii Rockwell array, with a pixel scale of $0''.148$ pixel $^{-1}$, giving a field of view of $\sim 150 \times 150$ arcsec 2 . The observations were performed in service mode under photometric conditions on 2010 March 6. The seeing, as derived from the FWHM size of stars in each frame, was mostly excellent during the observations, ranging from $0''.4$ to $0''.6$. Note that at the redshift of SPT 0538–50 observations in the *Ks*-band probe rest-frame are ~ 5800 \AA , or roughly the *R*-band.

The images were secured using individual exposures of two minutes per frame and a jitter procedure, which produces a set of frames at randomly offset telescope positions within a box of $10 \times 10''$. The total integration time was 30 minutes per target. Data reduction was performed using the ESO pipeline for jitter imaging data (Devillard 2001). Each frame was dark-subtracted and flat-fielded by a normalized flat field obtained from twilight sky images. Sky subtraction was performed using

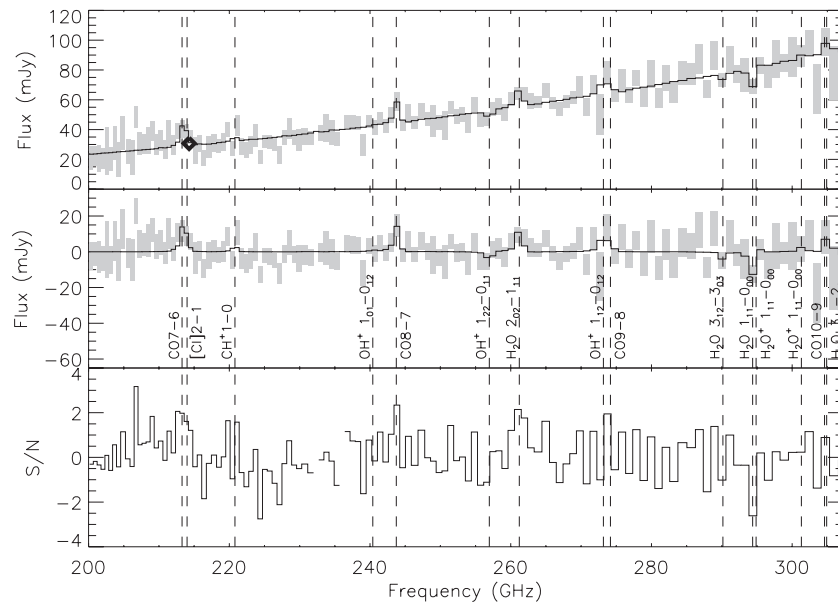


Figure 2. APEX/Z-Spec spectrum of SPT 0538–50. The top panel shows the spectrum with continuum, while the central panel shows the continuum-subtracted spectrum. The lower panel shows the S/N. The model spectrum (including all emission/absorption lines) is shown as a black overlay and all lines are identified and highlighted. See description in Section 2.2.1. The expected frequencies of other molecular lines lying within the band are also marked for reference. The redshift of SPT 0538–50 was originally obtained with this Z-Spec spectrum.

median averaged and scaled sky frames obtained combining jittered exposures of the same field. Sky-subtracted images were aligned to sub-pixel accuracy and co-added. Photometric calibration was performed from comparison with Two Micron All Sky Survey magnitudes of bright stars available in the field. The estimated internal photometric accuracy is ~ 0.1 mag.

2.1.8. *HST* WFC3-IR

SPT 0538–50 was observed with the *Hubble Space Telescope* (*HST*) Wide Field Camera 3 on 2011 October 10, as part of program 12659 (PI: Vieira). We use the F110W and F160W filters, the reddest high-throughput filters available. One complete orbit was dedicated to this source, split evenly between the two filters.

The data presented here are based on data products delivered by the *HST* pipeline, and photometry is defined following the WFC3 handbook.

2.2. Spectroscopy

Here we present observational details for the spectroscopic data used in this work. The VLT spectrum (Section 2.2.4) appeared previously in Hezaveh et al. (2013)—all other data are reported here for the first time.

2.2.1. Z-Spec Millimeter Spectroscopy

Z-Spec is a broadband, millimeter-wave direct-detection grating spectrometer operating from 185 to 305 GHz. The channel resolution, which is fixed by the spectrometer geometry, ranges from $R = 250$ to $R = 400$ (700 km s^{-1} to 1100 km s^{-1}). Further details about the Z-Spec instrument may be found in Naylor et al. (2003).

Observations of SPT 0538-50 were obtained from APEX on 2010 November 3 and 6 with a total of 352 minutes on-source.²⁹ Atmospheric conditions were excellent, with the opacity at

225 GHz in the range of 0.025–0.041. Calibration of each spectral channel was obtained from continuum observations of Mars, Uranus, and Neptune. The spectrum is shown in Figure 2, and is the source of the original redshift determination for SPT 0538–50. As a check on the calibration, we compared the SPT continuum flux at 1.4 mm with the Z-Spec flux and found them to be consistent.

Once the redshift was established (based on the detected $^{12}\text{CO}(8\rightarrow 7)$ and $^{12}\text{CO}(7\rightarrow 6)$ lines), we performed a fit of all likely species in the Z-Spec band with the redshift held fixed at $z = 2.7817$; the results are given in Table 1. For Z-Spec, the $^{12}\text{CO}(7\rightarrow 6)$ line is blended with the $[\text{C I}](2\rightarrow 1)$ due to the wide instrument channel width, and therefore the division between the line fluxes is uncertain. The fit places approximately equal flux in the two lines: $^{12}\text{CO}(7\rightarrow 6) = 16 \pm 8 \text{ Jy km s}^{-1}$ and $[\text{C I}](2\rightarrow 1) = 12 \pm 8 \text{ Jy km s}^{-1}$.

In addition to the CO and $[\text{C I}]$ lines, there are a number of H_2O transitions that are redshifted into this band. We identify two H_2O features; a 2.7σ detection of the $\text{H}_2\text{O}(2_{02} \rightarrow 1_{11})$ line in emission (as seen in several other galaxies: Bradford et al. 2009; Omont et al. 2011, 2013; Yang et al. 2013) and a 2.7σ detection of the $\text{H}_2\text{O}(1_{11} \rightarrow 0_{00})$ absorption line at 294 GHz, similar to the strength of that observed in Arp220 (Rangwala et al. 2011).

2.2.2. SMA

SPT 0538–50 was observed by the Submillimeter Array (SMA) in the compact configuration on 2010 November 9 and 27. The total on-source integration time was 2 hr. The upper and lower sidebands were centered around 225 and 213 GHz, respectively. Flux calibration was obtained by observing Callisto (November 9) and Uranus (November 27), while the quasar 3C84 was used for passband calibration. Complex gain calibration and astrometric verification were provided by the quasars J0538-440 and J0526-4830, respectively. Baselines ranged from 5–55 k λ , and the synthesized beam FWHM was $5'' \times 17''$. The redshifted $^{12}\text{CO}(7\rightarrow 6)$ and $[\text{C I}](^3P_2-^3P_1)$ lines (hereafter

²⁹ Part of ESO programs 086.A-0793 and 087.A-0815; PI: De Breuck; and 086.F-9318 and 087.F-9320; PI: Greve.

Table 1
All IR Emission Line Observations for SPT0538-50

Line	ν_{rest} (GHz)	ν_{obs} (GHz)	Flux (Jy km s ⁻¹)	L_{line} (10 ⁹ L _⊙)	$L_{\text{line}}/L_{\text{FIR}}$ (10 ⁻⁵)	Source
[O III]-52 μm	5786.0	1530.0	-25 ± 200	<70	<105	SPIRE-FTS
[N III]-57 μm	5230.6	1383.1	-20 ± 150	<70	<105	SPIRE-FTS
[O I]-63 μm	4744.9	1254.7	-160 ± 160	<75	<110	SPIRE-FTS
[O III]-88 μm	3393.1	897.2	150 ± 200	<50	<80	SPIRE-FTS
[N II]-122 μm	2459.4	650.3	30 ± 140	<40	<60	SPIRE-FTS
[O I]-146 μm	2060.1	544.8	-100 ± 180	<95	<145	SPIRE-FTS
[C II]-158 μm	1900.6	502.6	310 ± 190	<120	<185	SPIRE-FTS
¹² CO(10→9)	1151.9	304.6	Z-Spec
¹² CO(9→8)	1036.9	274.2	Z-Spec
¹² CO(8→7)	921.8	243.7	23 ± 10	3.1 ± 1.4	4.8 ± 2.1	Z-Spec
¹² CO(7→6)	806.7	213.3	16 ± 8	1.9 ± 0.9	2.9 ± 1.5	Z-Spec
¹² CO(7→6)	806.7	213.3	26.6 ± 4.5	3.5 ± 0.6	5.4 ± 1.0	SMA
¹² CO(1→0)	115.3	30.4	1.2 ± 0.2	0.02 ± 0.003	0.03 ± 0.006	ATCA
[C I](³ P ₂ → ³ P ₁)	809.3	214.0	12 ± 8	1.4 ± 0.9	2.2 ± 1.5	Z-Spec
[C I](³ P ₂ → ³ P ₁)	809.3	214.0	19.7 ± 2.6	2.3 ± 0.3	3.5 ± 0.6	SMA
H ₂ O(3 ₁₂ → 2 ₂₁)	1153.1	304.9	Z-Spec
H ₂ O(1 ₁₁ → 0 ₀₀)	1113.3	294.4	-19 ± 7	Z-Spec
H ₂ O(3 ₁₂ → 3 ₀₃)	1097.4	290.2	Z-Spec
H ₂ O(2 ₀₂ → 1 ₁₁)	987.9	261.2	19 ± 7	2.8 ± 1.0	4.3 ± 1.6	Z-Spec
H ₂ O(2 ₁₁ → 2 ₀₂)	752.0	198.8	Z-Spec

Notes. Undetected SPIRE lines have 2σ upper limits on their luminosities quoted. All values are as observed (i.e., they have not been corrected for the magnification factor of 21 ± 4). Fluxes for the undetected SPIRE lines have been measured by simply integrating at the expected line position, assuming a linewidth matching the ¹²CO(1→0) emission.

[C I](2→1)) were located in the lower sideband, while continuum information was derived from the upper sideband and line-free channels of the lower sideband. The correlator was configured for 3.25 MHz channels and a 4 GHz total bandwidth, corresponding to a 4.6 km s⁻¹ spectral resolution over the 5300 km s⁻¹ band.

The SMA data were reduced using the MIRIAD software package (Sault et al. 1995). Continuum images were created using multi-frequency synthesis to combine both sidebands, then CLEANed using natural weighting to 2 times the noise level to remove sidelobe effects. Spectral datacubes were produced by averaging velocity channels to a 90 km s⁻¹ resolution before inverting to the image plane.

The SMA 220 GHz spectrum is shown in Figure 3. The ¹²CO(7→6) and [C I](2→1) lines are each detected at $\sim 6\sigma$. We have shown a velocity-zeroed overlay of both detected emission lines (where zero velocity is defined at the targeted redshift, $z = 2.7817$, as discussed in Section 3.1).

2.2.3. Herschel SPIRE FTS

SPT 0538–50 was observed with the SPIRE Fourier Transform Spectrometer (FTS) on 2011 November 8 in single pointing mode as part of program OT1_dmarrone_1. Both observing bands were used: the short wavelength band, covering 194–313 μm (1544–959 GHz), and the long wavelength band, covering 303–671 μm (989–467 GHz). Observations were made in the high spectral resolution mode, resulting in a resolution of 0.04 cm⁻¹ (1.2 GHz) over each of the two bands. One hundred repetitions were carried out with a resultant on-source integration time of 13752 s.

The data were processed and calibrated using the standard SPIRE FTS pipeline. SPT 0538–50 is compact relative to the SPIRE FTS beam—our ALMA observations suggest the extent of the source is, at most, 5'' in extent in the far-IR, while the effective SPIRE beam varies from $\sim 17''$ to $\sim 42''$

across the spectrum. As such, we used calibration scripts appropriate for the reduction of a pure point source. For this relatively faint galaxy, the measured spectrum is dominated by emission from the instrument itself. This was removed using a combination of techniques, including subtraction of the average of adjacent spectral pixels and subtraction of a “dark sky” reference spectrum with the same integration time.

The SPIRE FTS spectrum of SPT 0538–50 is shown in Figure 4. A spectral baseline has been removed using a boxcar average over 30 channels. We fail to robustly detect any of the ISM lines that fall into the observed spectral range. We derive 2σ upper limits on line fluxes and luminosities following Valtchanov et al. (2011), by taking the expected line peak to be two times the rms noise (calculated within ± 5000 km s⁻¹ of the expected line centroid). Table 1 gives upper limits for the six ISM fine structure lines falling within the spectral range of our observations. We also give the value of the velocity-integrated flux at the position of each of the far-IR fine structure lines (where the velocity range has been chosen following the width of the ¹²CO(1→0) line).

There is a $\sim 2\sigma$ (relative to the local rms) excess of flux at the expected position of the [C II]-158 μm line, which can also be seen in a visual examination of the spectrum (see zoomed cutouts, Figure 4). Integrating the spectrum at the expected position of the [C II]-158 μm line, assuming a line-width that matches the ¹²CO(1→0) emission, results in a flux density of 310 ± 190 Jy km s⁻¹. Here, we treat [C II]-158 μm as a non-detection, and simply note the possible presence of a weak emission feature.

2.2.4. VLT XSHOOTER

SPT 0538–50 was observed by VLT/XSHOOTER on 2010 March 16 (PID 284.A-5029) for a total of 6×900 s integrations (see Vernet et al. 2011 for instrument details). Half of these were taken in poor weather, however, and the spectrum presented

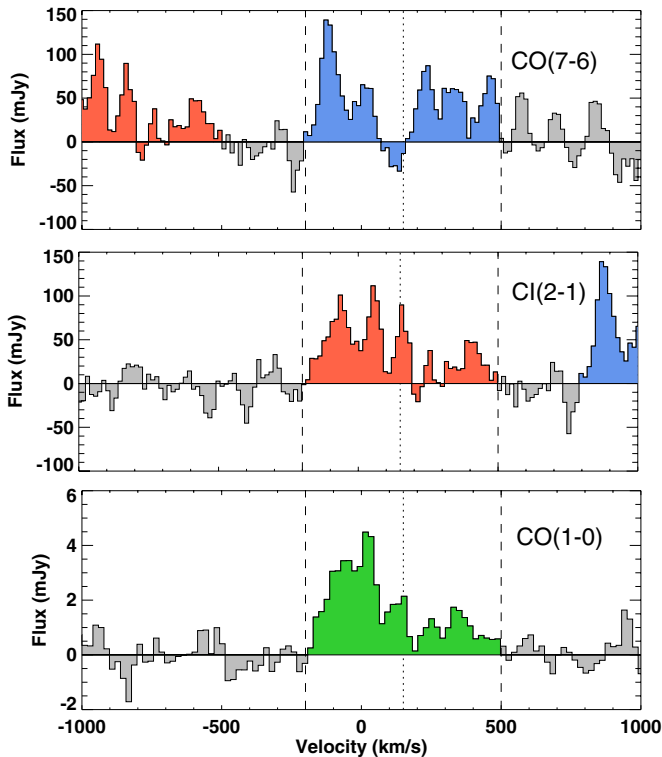


Figure 3. Spectra of molecular emission lines in SPT 0538–50, showing data from both the SMA and ATCA. All lines have been velocity-centered to a source redshift of $z = 2.7817$. Top panel: the $^{12}\text{CO}(7-6)$ emission line is detected at 213.37 GHz with the SMA and is highlighted in blue. Middle panel: the $[\text{C}\text{I}](2-1)$ line is also detected in the same SMA track and is shown in red. Note that this is the same SMA spectrum as shown in the top panel, but centered to a velocity appropriate for the $[\text{C}\text{I}](2-1)$ line. Lower panel: the velocity-centered ATCA 30 GHz spectrum, which detects the $^{12}\text{CO}(1\rightarrow 0)$ line, is shown in green. All lines show a double peaked structure (indicative of distinct kinematic components), and in all lines the blue peak is stronger than the red peak. Vertical dashed lines indicate the velocity range used to integrate line fluxes, and the vertical dotted line indicates the halfway point between the two.

(A color version of this figure is available in the online journal.)

here uses only three such integrations. The slit widths used were $1''$, $0.9''$, and $0.9''$ in the UV, visible, and NIR respectively. Further details of the observations are discussed in Hezaveh et al. (2013). Figure 5 shows the composite XSHOOTER spectrum of SPT 0538–50, with the wavelengths of lines (detected or undetected) in the source and lens labeled.

3. ANALYSIS

3.1. Redshift

The redshift of SPT 0538–50 is not usefully constrained by the millimeter-wavelength photometry from SPT alone. An initial redshift search was carried out using Z-Spec, using the technique described in Lupu et al. (2012) (which deals with low S/N lines by combining the significance of several detected lines). Several ISM emission lines were detected, including two high- J transitions of ^{12}CO : $^{12}\text{CO}(7\rightarrow 6)$ and $^{12}\text{CO}(8\rightarrow 7)$. Based on these lines, the redshift was determined to be $z = 2.783$.

This initial redshift was subsequently independently confirmed by the detection of the $^{12}\text{CO}(7\rightarrow 6)$ and $[\text{C}\text{I}]$ lines by the SMA and the strong detection of the $\text{CO}(1\rightarrow 0)$ line with ATCA (as discussed in Aravena et al. 2013), which both yielded a redshift of $z = 2.7817$.

We use optical spectroscopy to determine the redshift of the foreground lens, as reported in Hezaveh et al. 2013. The VLT/X-SHOOTER spectrum, shown in Figure 5, contains a multitude of absorption lines due to the foreground galaxy, and a redshift of $z = 0.441$ was determined by fitting Gaussian profiles to the Ca H+K absorption lines.

In addition to the foreground absorption lines, the ionized silicon line (Si iv at 1397 Å) can be identified as originating from SPT 0538–50, complete with a well-resolved P Cygni profile (characteristic of a rapid outflow). This identification would imply a redshift offset; the silicon line is detected blueward of the wavelength expected for $z = 2.7817$. The peak (i.e., of the red side) of the Si iv P Cygni profile is detected at 5427 Å, implying a line redshift of $z = 2.7692$. If interpreted as an outflow, this redshift offset corresponds to a velocity shift of $v \sim 1000 \text{ km s}^{-1}$.

This velocity offset can be compared to values in the literature for similar galaxies. Banerji et al. (2011) presents a correlation between outflow velocity and star formation rate for a variety of galaxy types. From their compilation, it can be seen that a star formation-driven outflow of $\sim 1000 \text{ km s}^{-1}$ is certainly not unusual for a galaxy such as SPT 0538–50 with a SFR of $760 M_{\odot} \text{ yr}^{-1}$ (see Section 3.4).

Other than Si iv, however, we do not detect any optical emission lines (such as Ly α or C iv 1550 Å) from SPT 0538–50. The non-detection of Ly α is unsurprising, due to the likely high column density of gas and dust present in a starbursting system such as SPT 0538–50.

3.2. NIR Morphology and Source/Lens De-blending

A major goal of our imaging program was to examine the morphological structure of SPT 0538–50. However, while observations at millimeter and submillimeter wavelengths provide an essentially unobscured view of the background DSFG, in the NIR (i.e., *HST*/WFC3 and *Spitzer*/IRAC) any emission from the background DSFG is strongly blended with that from the foreground lens. If we are to uncover the structure of the DSFG it is therefore necessary to de-blend the image, subtracting emission from the lens galaxy and leaving only the DSFG.

3.2.1. Modeling the Foreground Lens

We first create a model of the lens as observed by *HST*/WFC3 (using a co-add between the F110W and F160W bands). For this purpose, we chose the fitting code GALFIT 3 (Peng et al. 2010), which can simultaneously fit several source components and multiple neighboring galaxies (see Häußler et al. 2007)—vital for modeling a galaxy in a crowded field such as SPT 0538–50. If neighboring galaxies are not simultaneously modeled, the resulting fits can be poor.

We simultaneously fit the central elliptical galaxy and the nine closest neighbors. We attempted to fit the central galaxy with a standard two-component model, consisting of a central bulge plus an exponential disk. Fits with this two-component model were generally poor with unphysical parameters required in the model (convergence was only achievable with a bulge Sérsic index $n_b > 6$) and significant central structure remaining in the residuals. A three-component model, including a central point source, provided a superior fit with physically plausible parameters for each component and little residual structure. The effective radius of the “point source” component fitted by GALFIT is $750 \pm 20 \text{ pc}$. We allowed the Sérsic index of the bulge to vary as a free parameter with the result that the best-fitting model had a bulge Sérsic index of $n_b = 2.65$.

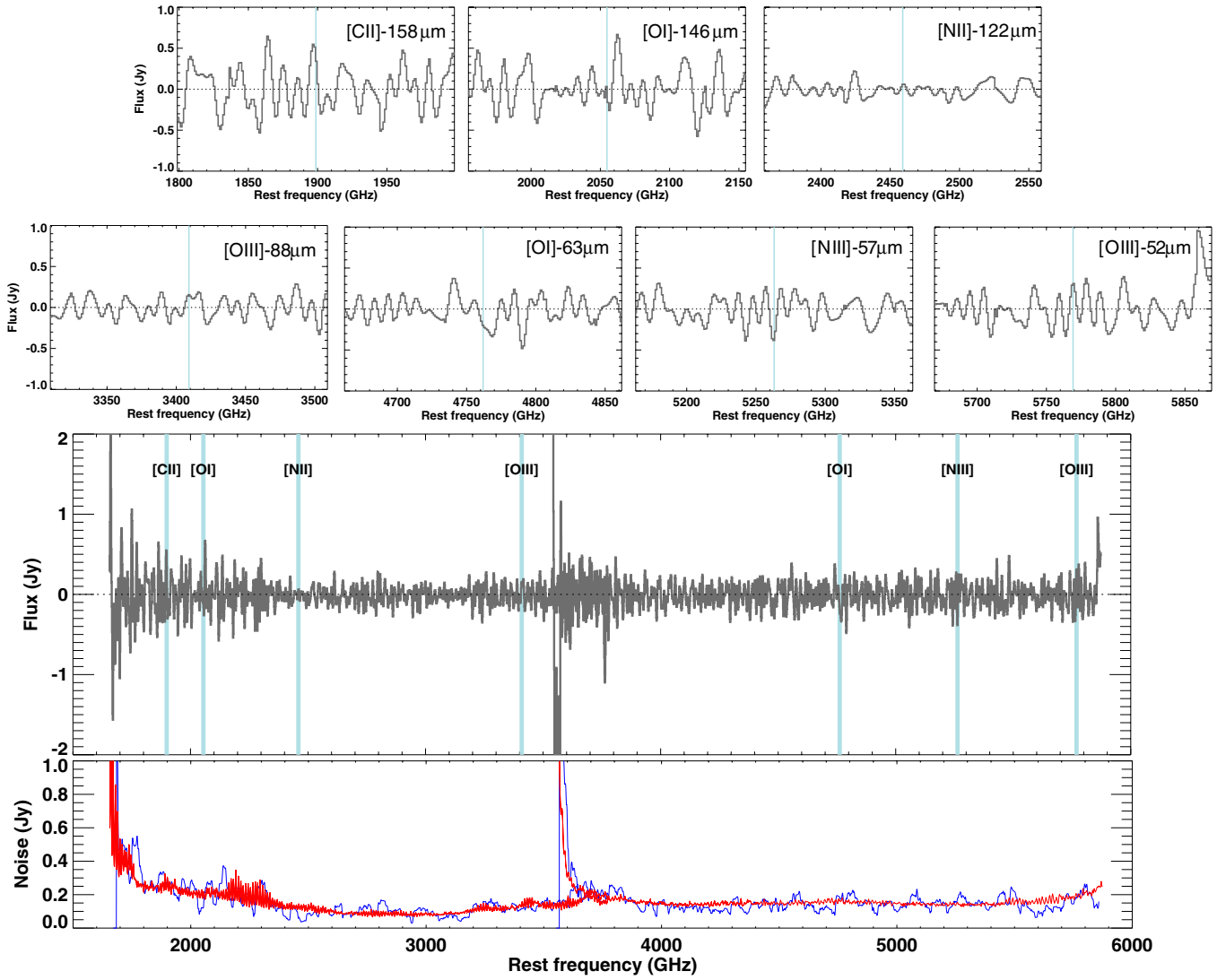


Figure 4. *Herschel* SPIRE FTS spectrum of SPT 0538–50. The center panel shows the spectrum of SPT 0538–50. The expected positions of some spectral lines are indicated. The upper panels show cutouts at the expected position of the lines zoomed to show ± 100 GHz around the line frequency. The lower panel shows the noise spectrum, both the uncertainty on the mean of the repeated scans (red) and the dispersion within ± 5000 km s $^{-1}$ of each channel. (A color version of this figure is available in the online journal.)

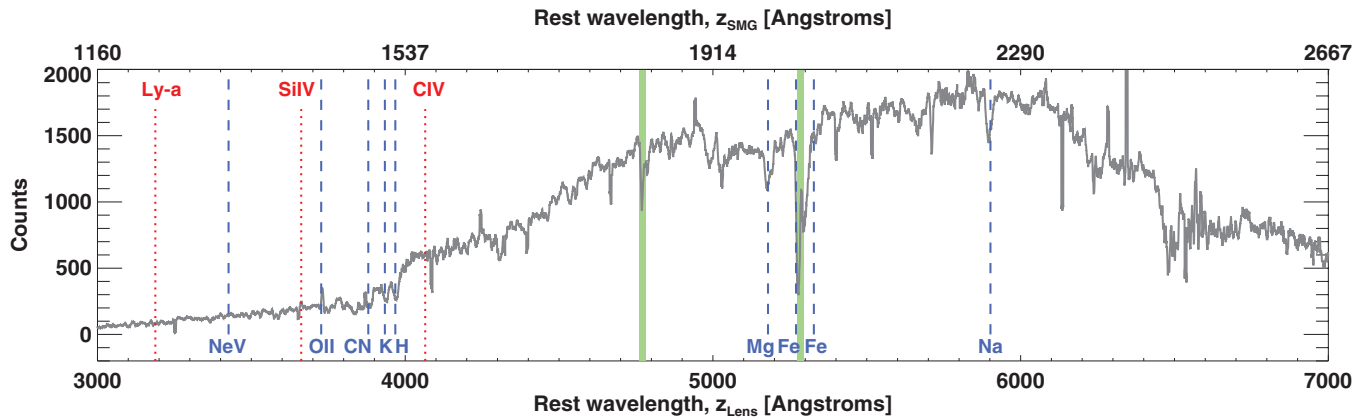


Figure 5. VLT/XShooter spectra of SPT 0538–50. The lower axis shows the rest wavelength at the redshift of the foreground lens ($z = 0.441$), while the upper axis shows the rest wavelength at the redshift of SPT 0538–50 ($z = 2.7817$). Spectral lines marked in blue are shown in the rest frame of the foreground lens, while lines marked in red are shown in the rest frame of the DSFG. The atmospheric A and B absorption bands are marked in green. (A color version of this figure is available in the online journal.)

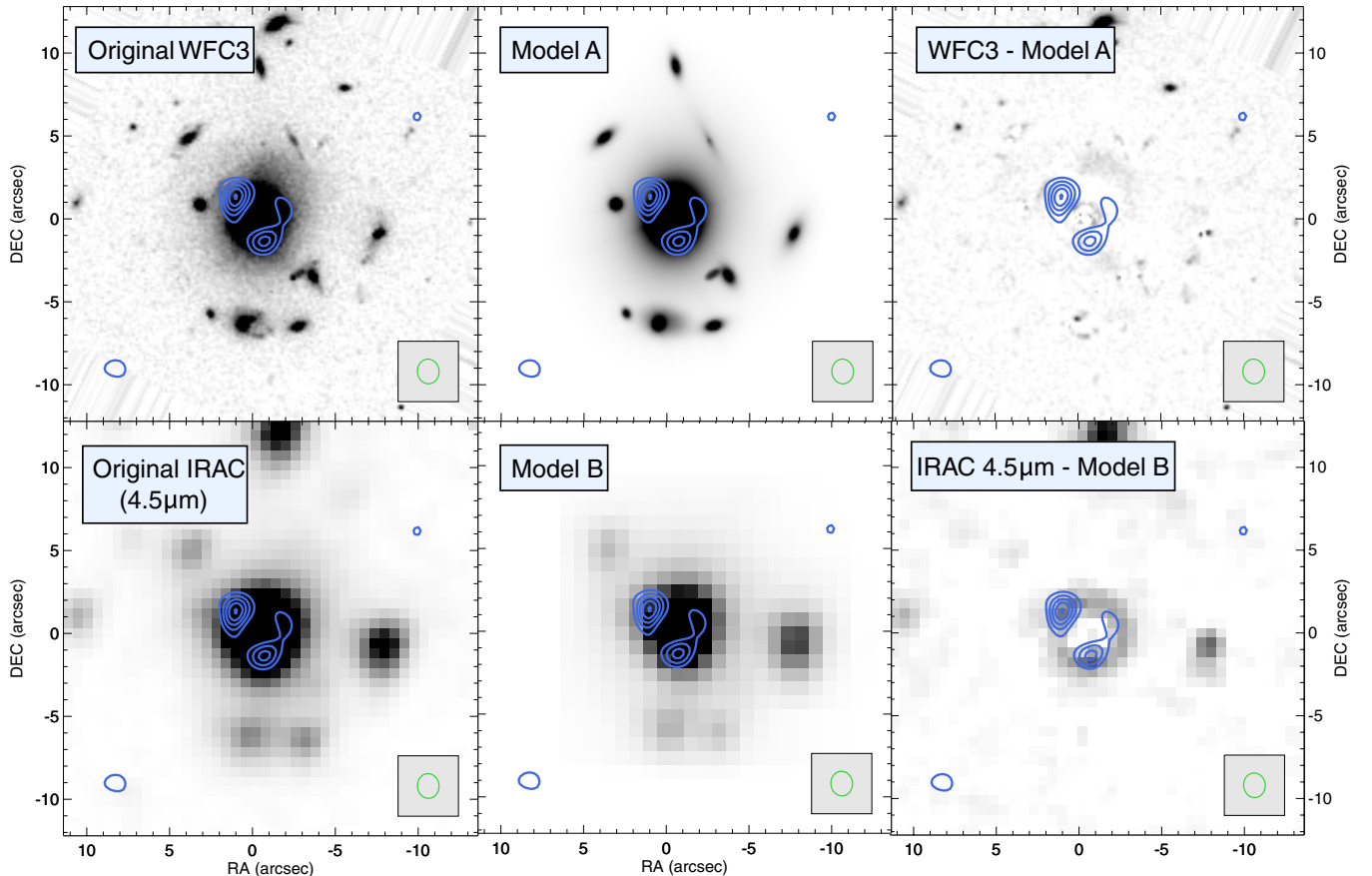


Figure 6. IR imaging of SPT 0538–50, showing our de-blending technique. In each $25'' \times 25''$ panel, blue contours show submillimeter continuum emission imaged by ALMA (see Section 2.1.2). Contours are in steps of 2σ , starting at 3σ . Top left: *HST*/WFC3 image of SPT 0538–50, created by co-adding images from the F110W and F160W filters. The high- z DSFG is totally hidden by the lens galaxy in the foreground. Top center: the best-fitting GALFIT model of SPT 0538–50 (“model A”). Top right: *HST*/WFC3 image of SPT 0538–50 after the subtraction of the best GALFIT model. The lower row shows our de-blending technique. Lower left: *Spitzer*/IRAC $4.5 \mu\text{m}$ image of SPT 0538–50. Lower center: The GALFIT model, convolved with the IRAC PSF at $4.5 \mu\text{m}$ (“model B”). Each object has been allowed to have a different *HST*–IRAC color. Lower right: *Spitzer*/IRAC $4.5 \mu\text{m}$ image of SPT 0538–50, after subtraction of the convolved GALFIT model. The Einstein ring structure of the DSFG, as traced by submillimeter continuum emission, is clearly visible in both IRAC channels once the foreground galaxy is removed. (A color version of this figure is available in the online journal.)

The top panels of Figure 6 show the original *HST*/WFC3 image (top left), the final “best-fit” model of the WFC3 data (top middle), and the residuals after subtracting this best-fit model from the WFC3 data (top right). We see no trace of the DSFG (the structure seen by ALMA) in these residuals; we thus conclude that the source detected in our *HST*/WFC3 imaging consists of flux from the low- z lens galaxy alone.

We note that the source as detected in IRAC (both $3.6 \mu\text{m}$ and $4.5 \mu\text{m}$) emission is extended on scales larger than those seen in the WFC3 image. The effective (i.e., half light) radius as measured in the IRAC images is $2''.0$ ($=11.3 \text{ kpc}$ at $z = 0.441$). In contrast, the effective radius in each of the WFC3 images is $1''.4$ ($=7.9 \text{ kpc}$ at $z = 0.441$). As the WFC3 residuals show no trace of the background SMG, this size difference further suggests the existence of blended emission from a background galaxy being detected in our IRAC images.

In addition, we also see no trace of emission from SPT 0538–50 remaining in the VLT/ISAAC K_s -band data. While this data has lower angular resolution than the *HST* images (and so is less suited to modeling the lens galaxy), we can use this K_s -band non-detection to place an upper limit on the flux from SPT 0538–50. Based on the rms noise in the ISAAC map, we place a K_s -band upper limit on SPT 0538–50 of $<20.19 \text{ mag}$.

3.2.2. De-blending the Source/Lens Emission

The next step is to construct a model of the lens galaxy as seen by IRAC. We take the high-resolution model of the lens galaxy developed from the *HST*/WFC3 imaging and convolve with the IRAC point-spread function (PSF). For this, we used an empirically derived PSF for which we used a uniformly weighted stack of ten bright, isolated stars in the same field. This model IRAC image can then be subtracted from the true IRAC image, leaving—as residual emission—any flux from the DSFG that was originally blended with the lens. The lower panels of Figure 6 show these steps—the original IRAC image (lower left), the above model convolved with the IRAC PSF (lower middle), and the residuals after subtracting this convolved model from the IRAC data (lower right).

It is of critical importance to determine the correct $3.6/4.5 \mu\text{m}$ normalization of the convolved galaxy model in order to recover lensed emission in the IRAC image. The normalization cannot be determined by making a direct comparison to the IR radial profile, or by the standard approach of minimizing the image residual, because the relatively large IRAC PSF will spread the contribution of any IR emission from SPT 0538–50 across much of the face of the lensing galaxy. To include the DSFG contribution to the image, we add a ring component at

the radial distance of the submillimeter components identified in the ALMA image, $1''.8$. It is important to note that this ring is a “dummy” component, and serves only as a guide to finding the optimum normalization of the model lens galaxy; it is not used in the final image subtraction, and so it does not contribute to any potential residuals.

The morphology of the DSFG is recovered once the properly normalized lens model image is removed. We varied the normalization of each component separately, using a simultaneous χ^2 fitting to find the optimum values determined by comparing them to the true IRAC image.³⁰

When convergence was reached, we subtracted the resulting best fitting model IRAC image from the real IRAC data. Having removed flux resulting from the foreground lens, we attribute the residuals left after this subtraction to the DSFG itself. We stress that ring-shaped residuals spatially matching the ALMA submillimeter emission are not artifacts, and are recovered even when using cruder normalization methods that do not use the dummy “ring” component. (See Appendix B for discussion and figures).

The results of this image deconvolution process are shown in Figure 6. ALMA submillimeter continuum contours have been overlaid on all images to indicate the position and structure of the lensed DSFG emission. The close spatial match between the ALMA continuum emission and the post-subtraction IRAC residuals strongly suggests that this residual NIR emission does indeed originate from the background DSFG.

We note that the poor subtraction of the galaxy $\sim 5''$ to the west of SPT 0538–50 is likely due to it being badly fit by a single Sérsic-profile. As it is significantly separated from SPT 0538–50, however, it will not affect our measured residual flux and we do not attempt to remove it more thoroughly.

We perform aperture photometry on the $3.6\ \mu\text{m}$ and $4.5\ \mu\text{m}$ residual images, extracting the flux within a circular aperture of radius $4''$ and derive resultant flux densities of $S_{3.6} = 22 \pm 5\ \mu\text{Jy}$ and $S_{4.5} = 47 \pm 8\ \mu\text{Jy}$. Uncertainties here include both the standard photometric uncertainty and the uncertainty resulting from the deconvolution process, as described above. In Section 4.3, we present these IRAC flux densities, de-magnified using our lens model.

The $[3.6]\ \mu\text{m}$ – $[4.5]\ \mu\text{m}$ color also provides evidence that the residual ring structure is a high- z background source. The lens galaxy has a $[3.6]$ – $[4.5]$ color of 0.43 ± 0.04 —fairly typical for an early-type galaxy at intermediate redshift. The residual structure, which we attribute to the background DSFG, has a far redder $[3.6]$ – $[4.5]$ color of 1.27 ± 0.5 , again suggesting that the residuals are not part of the lens galaxy, but have a separate, high- z , origin. A $[3.6]$ – $[4.5]$ color this red may be indicative of the presence of an active galactic nucleus (AGN) in SPT 0538–50 (see Yun et al. 2008), though with such large uncertainties this conclusion remains speculative.

3.3. Lens Model

SPT 0538–50 was modeled in detail by Hezaveh et al. (2013), based on our Cycle 0 ALMA Band 7 imaging. Details of the modeling technique can be found in that work. To summarize, the modeling is carried out in the u, v plane by modeling

³⁰ This differs somewhat from the approach of Hopwood et al. (2011), who used the peaks in their submillimeter images to locate Sérsic-profile DSFG components in the IR model. Here we make a more conservative assumption, namely that any lensed IR emission will appear somewhere on the ring that threads the submillimeter components, but not necessarily at the same positions. As such, our dummy ring component has equal flux at all points.

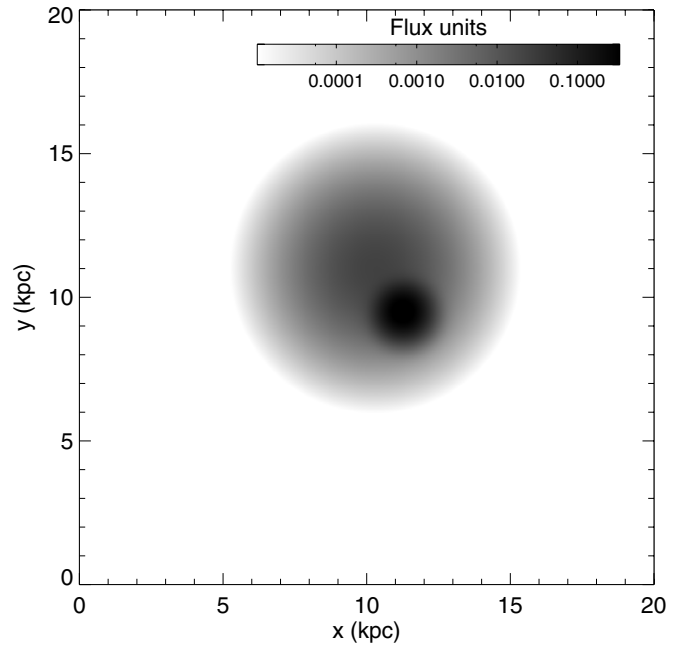


Figure 7. Source-plane reconstruction of SPT 0538–50 at $850\ \mu\text{m}$, based on the lens modeling reported in Hezaveh et al. (2013) and described in Section 3.3. Single-component fits to the lensed data were poor, with significant ($\sim 6\sigma$) residuals to the southeast. This two-component model provided a better fit to the data. The source consists of a large, diffuse component (responsible for 30% of the total luminosity) and a compact, luminous component (responsible for the remaining 70%).

the source as a two-dimensional Gaussian intensity profile, predicting u, v visibilities, and minimizing an χ^2 figure of merit.

As discussed by Hezaveh et al. (2013), the large size of the Einstein ring resolved by ALMA (diameter $\sim 1''.8$; see Figure 1) suggests a particularly massive lensing halo, with a mass of $(7.15 \pm 0.05) \times 10^{11} M_{\odot}$ enclosed within the Einstein radius. This massive foreground galaxy, combined with an optimal source/caustic configuration, results in a large magnification factor, $\mu = 21 \pm 4$. Fits using just a single component model in the source plane resulted in a significant ($> 6\sigma$) residual structure to the southeast of the image, evidence that a two-component model is required to reproduce the ALMA data. The best fitting model of SPT 0538–50 consists of two bodies, one compact ($R_{1/2} = 0.52 \pm 0.12\ \text{kpc}$) source, which is responsible for 70% of the total source-plane luminosity, and one more extended ($R_{1/2} = 1.61 \pm 0.33\ \text{kpc}$) source, which is responsible for the remaining 30%. Figure 7 shows the source-plane reconstruction of SPT 0538–50.

Using the lens model, it is now possible to begin reconstructing the intrinsic properties of SPT 0538–50. It can be seen that de-lensing its prodigious observed submillimeter flux density, $125 \pm 7\ \text{mJy}$ at $870\ \mu\text{m}$, results in a flux density of $6.1 \pm 1.2\ \text{mJy}$,³¹ which is relatively modest for a submillimeter galaxy. It is worth noting that the intrinsic submillimeter flux density still falls above the formal selection limit for the original SCUBA survey (Smail et al. 1997), and so this source would have been identified as a “classical” DSFG even if unlensed.

3.4. SED Fitting to IR Data

We use the galaxy SED fitting code CIGALE (Burgarella et al. 2005; Noll et al. 2009) to fit the near- and far-IR photometric

³¹ Uncertainties on this value, and on all de-lensed quantities quoted hereafter, include the uncertainty contribution from the lensing model.

Table 2
IR and Millimeter Fluxes for SPT 0538–50, Both Observed and De-magnified by the Lensing Magnification of 21 ± 4

Wavelength (Observed) λ	Observed Flux Density (mJy)	Innate Flux Density (mJy)	Source	Reference
1.1 μm	$<2.4 \times 10^{-4}$	$<5.0 \times 10^{-3}$	<i>HST</i> /WFC3	This work
1.6 μm	$<4.0 \times 10^{-4}$	$<8.4 \times 10^{-3}$	<i>HST</i> /WFC3	This work
2.2 μm	$<2.6 \times 10^{-4}$	$<5.5 \times 10^{-3}$	<i>VLT</i> /ISAAC	This work
3.6 μm	0.022 ± 0.005	0.0011 ± 0.0005	<i>Spitzer</i> /IRAC	This work
4.5 μm	0.047 ± 0.008	0.0023 ± 0.0006	<i>Spitzer</i> /IRAC	This work
100 μm	31 ± 2	1.5 ± 0.3	<i>Herschel</i> /PACS	This work
160 μm	141 ± 15	6.7 ± 1.5	<i>Herschel</i> /PACS	This work
250 μm	326 ± 23	15.5 ± 3.1	<i>Herschel</i> /SPIRE	This work
350 μm	396 ± 38	18.9 ± 3.9	<i>Herschel</i> /SPIRE	This work
500 μm	325 ± 24	15.5 ± 3.2	<i>Herschel</i> /SPIRE	This work
350 μm	336 ± 88	16.4 ± 5.1	APEX/SABOCA	Greve et al. (2012)
870 μm	125 ± 7	6.1 ± 1.2	APEX/LABOCA	Greve et al. (2012)
1.4 mm	28 ± 4.6	1.5 ± 0.4	SPT	Vieira et al. (2010)
2 mm	8.5 ± 1.4	0.4 ± 0.1	SPT	Vieira et al. (2010)
8.6 mm	0.13 ± 0.02	0.006 ± 0.002	ATCA	Aravena et al. (2013)

Notes. Quoted uncertainties include photometric flux uncertainties, magnification uncertainty (where applicable), and absolute flux calibration uncertainties.

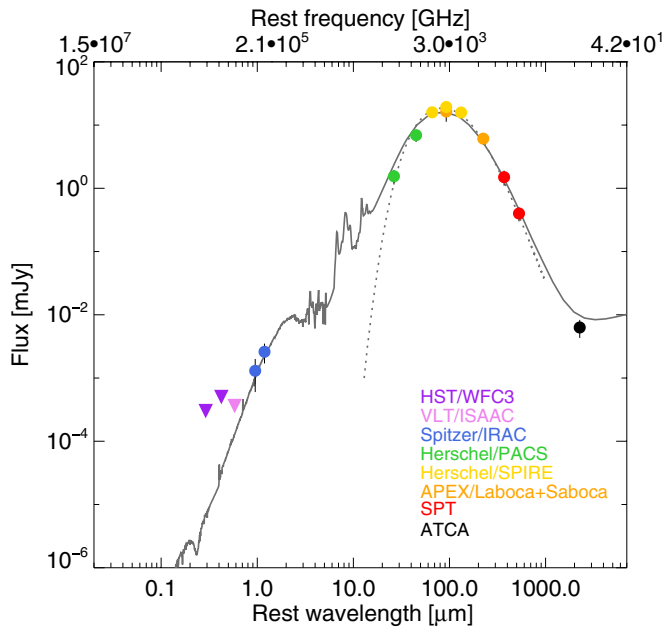


Figure 8. Result of SED fitting our SPT 0538–50 photometry, using the SED fitting code CIGALE. Fluxes are innate (i.e., demagnified using our lens model). Datapoints are ATCA 35 GHz continuum, SPT 2.0+1.4 mm, APEX 350 μm +870 μm , *Herschel*/SPIRE 250 μm +350 μm +500 μm , *Herschel*/PACS 100 μm +160 μm , *Spitzer*/IRAC 3.5 μm +4.5 μm , and 2σ upper limits from *HST*/WFC3 and *VLT*/ISAAC. The dashed line shows a graybody fit to the far-IR data with a dust temperature of 39 ± 2 K (and $\beta = 2$).

(A color version of this figure is available in the online journal.)

data. Table 2 gives de-magnified fluxes for all our observed IR datapoints, which sample both the stellar mass and the dust component of SPT 0538–50. We do not fit to the low frequency (<10 GHz) ATCA data (given in Aravena et al. 2013), as CIGALE does not extend its fitting beyond the far-IR.

The best-fitting SED, along with the de-magnified IR datapoints, is shown in Figure 8. The fit was achieved assuming a solar metallicity, an old stellar population with a mean age of 2 Gyr, and a young stellar population with a mean age of 20 Myr (these three input parameters are fixed). We also assume that the

older stellar population has an exponentially declining star formation history (SFH), while the young stellar population has a “box”-like SFH.

From these fixed input parameters, we then used CIGALE to find the best-fitting SFR, stellar mass, and IR luminosity. Being constrained by the FIR region of the SED, the SFR of SPT 0538–50 remains approximately constant when varying our input parameters. The stellar region (in the rest-frame optical/UV) is less well constrained, with only the two IRAC data points available. As a result, the derived stellar mass is more sensitive to the choice of input parameters. The stellar mass is only weakly dependent on the SFH assumed but is more dependent on the assumed age for the young stellar population. SED fits with ages <15 Myr were unable to match the derived IRAC fluxes. Varying the adopted young stellar age between 15–50 Myr causes the resulting stellar mass to vary by $\sim 50\%$. The stellar mass is also dependent on the assumed “burst fraction”—the mass fraction of the young stellar population. We assumed a fixed burst fraction of 0.3 (as with the stellar ages above, burst fractions of <0.3 provided unsatisfactory fits). It is difficult to constrain the burst fraction without more extensive optical photometry; however, allowing the burst fraction to vary within a reasonable range (0.3–0.6) alters the resulting stellar mass by less than 25%.

The SED suggests an IR luminosity (from 8–1000 μm) for SPT 0538–50 of $(3.4 \pm 0.3) \times 10^{12} L_{\odot}$, and an associated SFR of $760 \pm 100 M_{\odot} \text{ yr}^{-1}$. These values are in line with those calculated by fitting a simple graybody to the far-IR points alone; once corrected for the lensing factor used in this work, the IR luminosity of SPT 0538–50 derived by Greve et al. (2012) is $(3.17 \pm 1.0) \times 10^{12} L_{\odot}$. The SED fit also suggests a stellar mass of $(3.3 \pm 1.5) \times 10^{10} M_{\odot}$. We note that this is roughly consistent with the value of $(4.6 \pm 2.1) \times 10^{10} M_{\odot}$ reported in Aravena et al. (2013), which was estimated using an empirical relation between the specific star formation rate (sSFR = SFR/ M^*) and radio spectral flattening as observed for local compact starbursts (see Murphy et al. 2013). From these results, we also derive an sSFR for SPT 0538–50 of $25 \pm 12 \text{ Gyr}^{-1}$. The properties of SPT0538-50 derived from this analysis are summarized in Table 3.

Table 3
Derived Physical Properties for SPT 0538–50

Parameter	Value
SFR	$760 \pm 100 M_{\odot} \text{ yr}^{-1}$
M_{*}	$(3.3 \pm 1.5) \times 10^{10} M_{\odot}$
sSFR	$25 \pm 12 \text{ Gyr}^{-1}$
M_{H_2}	$(1.7 \pm 0.5) \times 10^{10} M_{\odot}$
L_{IR}	$(3.4 \pm 0.3) \times 10^{12} L_{\odot}$
$\Sigma_{\text{SFR}} \text{ (compact)}$	$(630 \pm 240) M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$
$\Sigma_{\text{SFR}} \text{ (diffuse)}$	$(31 \pm 11) M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$

4. PHYSICAL PROPERTIES

4.1. Source Structure

As described above, SPT 0538–50 is best modeled as being comprised of two source-plane components: a luminous compact body with $R_{1/2} = 0.52 \pm 0.12$ kpc and a more diffuse body with $R_{1/2} = 1.61 \pm 0.33$ kpc (models assuming just a single component were a poor fit to the data, resulting in $>6\sigma$ residuals). We note that the lensing models here assume symmetric Gaussian profiles for the source components, and as such do not take into account the likely complexity of the true galaxy.

Being a luminous DSFG, SPT 0538–50 is potentially undergoing a major merger event that is driving the high rate of star formation (e.g., Engel et al. 2010). However, it is not immediately obvious whether the two components revealed by the lensing reconstruction correspond to separate bodies involved in a merger or whether the compact, luminous component simply represents a site of star formation embedded within a more diffuse ISM.

Taking the two source-plane components of SPT 0538–50 separately, we use the projected radii and luminosities to calculate the SFR density in each of the two components of SPT 0538–50. By assuming that the flux density distribution revealed in the ALMA imaging (70/30 in favor of the compact component) can be applied to the total IR luminosity, we calculate that the SFR density (Σ_{SFR}) of the compact component is over an order of magnitude higher than its diffuse companion. For the compact and extended bodies, we reach values of $\Sigma_{\text{SFR}} = 630 \pm 240 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$ and $\Sigma_{\text{SFR}} = 31 \pm 11 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$, respectively.

The majority (70%) of the total star formation in SPT 0538–50, therefore, likely occurs in a compact region that has an SFR surface density comparable to the densest star-forming regions contained within local merging ULIRGs (i.e., Kennicutt 1998; Genzel et al. 2010). This is suggestive of a galaxy undergoing a major merger (in which tidal torques funnel molecular material into a dense, star-forming region) rather than secular star formation (which is expected to produce a more extended, disk-like system). The fact that the majority of the star formation in SPT 0538–50 is occurring in a dense, ULIRG-like environment is supported by both the highly excited CO emission discussed in Section 4.2, and the far-IR line deficits discussed in Section 4.3.

These conclusions are further supported by the derived sSFR of SPT 0538–50 of $\sim 25 \text{ Gyr}^{-1}$ (see Section 3.4). This value is approximately a factor of five above the “main sequence” at $z = 2.7$ (e.g., Elbaz et al. 2011; Karim et al. 2011), which again suggests the presence of a major-merger driven starburst. The double-peaked molecular lines revealed by our millimeter-wavelength spectroscopy are certainly in line with this picture. However, in the absence of high-resolution

kinematic observations, it is difficult to draw firm conclusions about the source structure.

4.2. ISM Properties

The velocity-resolved measurements of CO and C_I (Figure 3) reveal SPT 0538–50 to be a double-peaked source. Emission from the low- J CO and C_I lines is expected to be co-spatial, as the two molecules have similar critical densities and excitation temperatures, to the extent that C_I is also expected to be a good tracer of the total molecular gas mass (Papadopoulos & Greve 2004).

As our SMA-detected molecular lines do not have particularly Gaussian profiles, we determine line fluxes by integrating our spectra between velocity limits (-200 to $+500 \text{ km s}^{-1}$) defined by the profile of the CO(1–0) emission line detected with high S/N in our ATCA spectrum (see Section 2.1.3 and Figure 3). We find the $^{12}\text{CO}(7-6)$ line has a flux of $S_{\text{CO}(7-6)} = 26.6 \pm 4.5 \text{ Jy km s}^{-1}$, while the $[\text{C I}](^3P_2-^3P_1)$ has a flux $S_{[\text{C I}]} = 19.7 \pm 3.1 \text{ Jy km s}^{-1}$.

For both the CO(1–0) and C_I lines the “red” peak is suppressed relative to the “blue” peak, though this effect is less evident in the CO(7–6) line. We can integrate the peaks separately (above and below 150 km s^{-1} , the point that lies between the two peaks) to assess the strength of the red peak suppression. The CO(1–0) and C_I lines each have $\sim 80\%$ of their line flux in their blue peak, while the CO(7–6) line shows far more symmetry between the peaks (with the blue peak containing $\sim 60\%$ of the total line flux). The asymmetry between the red and blue peaks seen in the CO(1–0) and C_I lines suggests that they represent distinct components of an on-going major merger, as opposed to being due to the rotation of a disk.

4.2.1. Molecular Gas Properties

As discussed in Section 2.1.2, SPT 0538–50 has been observed by ATCA in $^{12}\text{CO}(1\rightarrow 0)$ (Aravena et al. 2013). They report a de-lensed³² $^{12}\text{CO}(1\rightarrow 0)$ flux of $I_{\text{CO}(1-0)} = 1.20 \pm 0.20 \text{ Jy km s}^{-1}$, and a resulting $^{12}\text{CO}(1\rightarrow 0)$ luminosity of $L'_{\text{CO}(1-0)} = (2.16 \pm 0.13) \times 10^{10} \text{ K km s}^{-1} \text{ pc}^2$. Using a standard ULIRG-appropriate CO/H₂ conversion factor,³³ this corresponds to a molecular gas mass of $1.7 \pm 0.5 \times 10^{10} M_{\odot}$ (the uncertainty of this value incorporates both the uncertainty of the lensing magnification and an estimated 15% calibration uncertainty). This is toward the lower end of the molecular gas mass distribution for DSFGs (Bothwell et al. 2013 report a mean molecular gas mass for unlensed DSFGs of $5.3 \pm 1.0 \times 10^{10} M_{\odot}$), but is not incommensurate with the derived SFR and stellar mass. We calculate the IR-to-H₂ ratio (defined here as $\text{SFR}/M(\text{H}_2)$) of $4.4 \pm 1.4 \times 10^{-8} \text{ yr}^{-1}$, and a baryonic gas fraction—defined as $M(\text{H}_2)/(M(*)+M(\text{H}_2))$ —of $f_{\text{gas}} \sim 0.4$.

Aravena et al. (2013) used the peak separation of the $^{12}\text{CO}(1\rightarrow 0)$ line to estimate a dynamical mass for SPT 0538–50. Taking the source separation (2 kpc) and the $^{12}\text{CO}(1\rightarrow 0)$ peak separation above, they estimate a dynamical mass for SPT 0538–50 of $\sim 1.3 \times 10^{11} M_{\odot}$.

As the $^{12}\text{CO}(1\rightarrow 0)$ line is a tracer of the cold molecular gas reservoir and the $^{12}\text{CO}(7\rightarrow 6)$ line traces warmer star-forming gas, we can use the ratio of the $^{12}\text{CO}(7\rightarrow 6)$ and $^{12}\text{CO}(1\rightarrow 0)$

³² The de-lensing was done using the magnification factor of $\mu = 21 \pm 4$ used throughout this work.

³³ This is itself uncertain—the standard value for starburst galaxies, which we adopt here, is $\alpha_{\text{CO}} = 0.8 M_{\odot} (\text{K km s}^{-1} \text{ pc}^2)^{-1}$ (see Solomon & Vanden Bout 2005; Carilli & Walter 2013).

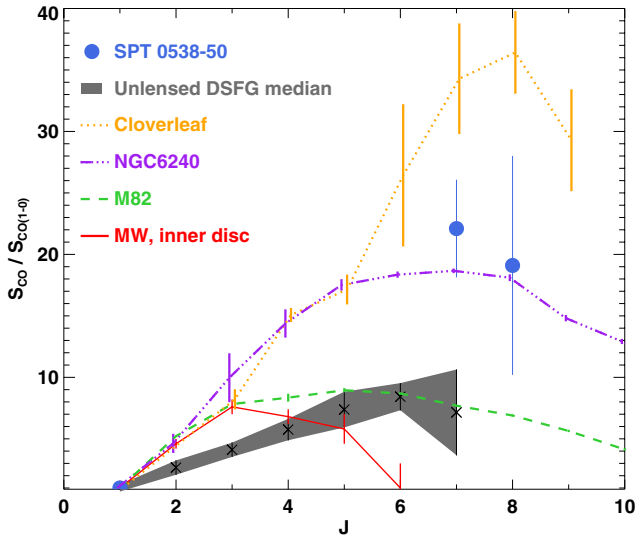


Figure 9. CO spectral line energy distribution (SLED) normalized to the $^{12}\text{CO}(1\rightarrow 0)$ flux for SPT 0538–50. Also shown is the median SLED for a large sample of unlensed DSFGs (Bothwell et al. 2013); the SLED for local LIRG NGC 6240 (Meijerink et al. 2013), the prototypical low- z starburst M82 (N. Rangwala 2013, private communication—see also Panuzzo et al. 2010; Kamenetzky et al. 2012), the inner disk of the Milky Way (Fixsen et al. 1999), and the lensed Cloverleaf QSO (Bradford et al. 2009; Riechers et al. 2011). SPT 0538–50 shows a high excitation toward higher J relative to similar starburst galaxies.

(A color version of this figure is available in the online journal.)

luminosities as a crude estimator of the ISM excitation in SPT 0538–50. Figure 9 shows the CO SLED of SPT 0538–50 compared to SLEDs of similar objects, including the “Eyelash” DSFG as presented by Danielson et al. (2011) and the median SLED for a sample of unlensed DSFGs (Bothwell et al. 2013). It can be seen that the CO line flux from SPT 0538–50 in our two high- J lines, $(8\rightarrow 7)$ and $(7\rightarrow 6)$, is far stronger than would be extrapolated from the $(1\rightarrow 0)$ line alone (assuming a CO SLED comparable to other DSFGs). We find a ratio of $L'_{\text{CO}(7-6)}/L'_{\text{CO}(1-0)} = 0.44 \pm 0.08$, compared to 0.12 ± 0.005 for the “Eyelash” (Danielson et al. 2011) and 0.18 ± 0.04 for the DSFG average (Bothwell et al. 2013).

In the absence of further detections of mid- J lines (i.e., $2 < J < 6$) to constrain the behavior of the SLED, it is difficult to draw firm conclusions about the excitation conditions of the ISM in SPT 0538–50. It is possible that the strong CO emission at high J indicates a star-forming ISM, which displays a higher excitation than observed in other, star formation-dominated systems. A possible AGN could also contribute to the observed CO excitation, though this is somewhat unlikely as the bulk of the submillimeter-selected DSFG population is likely to be starburst-dominated (Coppin et al. 2010).

Alternatively, the extreme excitation may also result from differential lensing of the CO transitions. As discussed by Serjeant (2012) and Hezaveh et al. (2012), compact regions can be magnified more strongly than the diffuse ISM as a whole by lying close to a lensing caustic. If the compact region being boosted is a star-forming “clump” (with a high SFR density and high excitation), the effect on the SLED of SPT 0538–50 will be to produce an apparently higher excitation by disproportionately enhancing the contribution from a compact, bright region.

4.2.2. [C I]

We can use our measurement of the $[\text{C I}](2\rightarrow 1)$ line to estimate the total mass of carbon in SPT 0538–50. Following

Walter et al. (2011), who surveyed atomic carbon emission in 10 DSFGs, we note that the typical $[\text{C I}](1\rightarrow 0)/[\text{C I}](2\rightarrow 1)$ ratio (excluding sources with only upper limits on both lines) is 0.45 ± 0.14 , corresponding to an excitation temperature of 29.1 K. Assuming this value and the observed line flux, the resulting carbon mass is $(5.6 \pm 1.7) \times 10^6 M_{\odot}$, comparable to measured values for other DSFGs (Walter et al. 2011), and the carbon abundance is $M(\text{C I})/M(\text{H}_2) = (3.3 \pm 1.2) \times 10^{-4}$. We also derive the cooling contribution of C I (parameterized as $L_{[\text{C I}]} / L_{\text{FIR}}$, finding a value of $(9.7 \pm 3.6) \times 10^{-6}$. This value is again similar to other DSFGs (Walter et al. 2011), but higher than what is observed in QSOs at similar redshifts (Carilli & Walter 2013).

4.2.3. Water Detection and Excitation Conditions

Our Z-Spec observations also tentatively detect an emission line of water, $\text{H}_2\text{O}(2_{02} \rightarrow 1_{11})$ at 261 GHz. With a flux of $19 \pm 7 \text{ Jy km s}^{-1}$, this line is bright with a (de-lensed) luminosity of $1.4 \times 10^8 L_{\odot}$. The water emission detected from SPT 0538–50 is similar to that in other high- z systems. Omont et al. (2013) report the detection of water in six lensed ULIRGs lying between $1.5 < z < 4$, finding that water emission lines typically exhibit flux strengths $\sim 30\%$ – 50% of the nearby CO emission lines. In SPT 0538–50 the $\text{H}_2\text{O}(2_{02} \rightarrow 1_{11})$ line flux is $80\% \pm 30\%$ of the neighboring $\text{CO}(8\rightarrow 7)$ line.

Yang et al. (2013) present observations of various transmissions of water for a large sample of local and high- z galaxies (including lensed ULIRGs at high z). All transitions of water show a good correlation with IR luminosity—for their sample of galaxies, the $\text{H}_2\text{O}(2_{02} \rightarrow 1_{11})$ line has a mean $L(\text{H}_2\text{O})/L(\text{IR})$ ratio of 7.58×10^{-6} , with low scatter. The water emission from SPT 0538–50 is comparatively over-luminous: the $L(\text{H}_2\text{O})/L(\text{IR})$ ratio, again measured using the $\text{H}_2\text{O}(2_{02} \rightarrow 1_{11})$ line, is $(4.3 \pm 1.6) \times 10^{-5}$, over a factor of five times brighter than the mean value shown by the Yang et al. (2013) sample. SPT 0538–50 has the highest $L(\text{H}_2\text{O})/L(\text{IR})$ ratio even amongst the lensed high- z ULIRGs (which themselves show water emission more luminous than average).

This water emission, combined with the likely high level of excitation present in SPT 0538–50, is again consistent with the presence of a compact and extreme star-forming region (see González-Alfonso et al. 2010), though—as with the ^{12}CO emission discussed above—the observed water excitation may be enhanced by a differential lensing effect.

4.3. Far-IR Fine Structure Lines

Section 2.2.3 describes our *Herschel*/SPIRE FTS observations of SPT 0538–50. As discussed above, we fail to robustly detect any of the far-IR fine structure lines falling within the band. Figure 10 shows our derived upper limits (in units of $L_{\text{line}}/L_{\text{FIR}}$), compared to the typical range of values for starburst systems compiled by Spinoglio et al. (2012). With the possible exception of $[\text{O I}]63 \mu\text{m}$ (see below), all our upper limits are consistent with SPT 0538–50 having line luminosity ratios comparable to other starburst and IR-luminous galaxies.

It has long been known that the ratio between far-IR line luminosity and far-IR continuum emission has a connection to the total IR luminosity of a galaxy, with luminous and ultra-luminous IR galaxies (LIRGs and ULIRGs) exhibiting “deficits” in their far-IR fine structure line ratios relative to the values for secular, non-starburst galaxies (e.g., Luhman et al. 2003; Abel et al. 2009). One explanation for this result is the effect of the ionization parameter U (defined as the

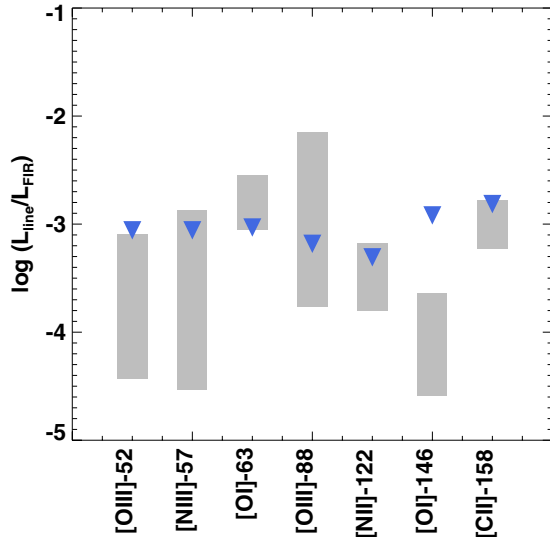


Figure 10. Comparison of the FIR line upper limits with typical values for local (U)LIRGs. Blue triangles show the 2σ upper limits on the FIR fine structure lines in SPT 0538–50, derived using the SPIRE spectrum. Gray regions show the 1σ range in values for (U)LIRG galaxies compiled by Spinoglio et al. (2012). (A color version of this figure is available in the online journal.)

ratio between hydrogen ionizing photons and hydrogen atoms; Abel et al. 2005). In this model, galaxies with enhanced ionization parameters have a greater fraction of their UV photons absorbed by interstellar dust and thus fewer photons available to photoionize the gas (e.g., Abel et al. 2009).

As discussed by Graciá-Carpio et al. (2011), a good proxy for the ionization parameter U is the ratio between far-IR luminosity and molecular gas mass, $L_{\text{FIR}}/M(\text{H}_2)$ —galaxies with enhanced values of $L_{\text{FIR}}/M(\text{H}_2)$ tend to have weaker far-IR fine structure lines, with a “transition” toward weaker lines happening at $>80 L_{\odot} M_{\odot}^{-1}$. We calculate SPT 0538–50 to have an $L_{\text{FIR}}/M(\text{H}_2)$ ratio of $(250 \pm 70) L_{\odot} M_{\odot}^{-1}$. This value would number among the highest recorded in the compilation by Graciá-Carpio et al. (2011), corresponding to the galaxies with the weakest far-IR fine structure lines. This value is high even for luminous DSFGs—the compilation of 32 CO-observed galaxies presented in Bothwell et al. (2013) has a mean (\pm standard deviation) $L_{\text{FIR}}/M(\text{H}_2)$ ratio of $(115 \pm 80) L_{\odot} M_{\odot}^{-1}$. The extreme $L_{\text{FIR}}/M(\text{H}_2)$ ratio observed in SPT 0538–50 suggests an unusually high ionization parameter, which would result in extremely weak far-IR fine structure lines, consistent with the lack of detection of lines in our SPIRE spectrum.

Recently, some results have suggested that the [O I]63 μm line ratio may be higher in some high- z DSFGs than in local (U)LIRGs—that is, the DSFGs may exhibit less of a [O I]63 μm line deficit than nearby merging systems. Coppin et al. (2012) found that a sample of DSFGs in the LABOCA Extended Chandra Deep Field South exhibit [O I]63 μm to far-IR ratios of approximately 0.5–1.5 percent—an order of magnitude brighter than seen in local (U)LIRGs. Those authors argue that the relatively “enhanced” [O I]63 μm line strength suggests that the star formation mechanism in their sample of DSFGs is more comparable to secular, disk galaxies (which typically do not display a [O I]63 μm deficit).

Our upper limit on the [O I]63 μm line luminosity in SPT 0538–50 is sufficient to constrain the [O I]63 μm to far-IR ratio to be significantly below the values recorded by Coppin et al. (2012), and at the lower end of the range for local

(U)LIRGs. According to the model described above, this would suggest that star formation is more compact in SPT 0538–50 than in the unlensed DSFGs in the Coppin et al. (2012) sample.

5. CONCLUSIONS

We have presented a detailed observational study of a $z \sim 2.8$ DSFG, which is being strongly gravitationally lensed by an intermediate-redshift elliptical galaxy. The lensing magnification allows us to carry out an array of multi-wavelength observations. A detailed lens model constructed from ALMA Band 7 data (Hezaveh et al. 2013) reveals that SPT 0538–50 consists of two source-plane components (a bright compact component, and a dimmer diffuse component), and is magnified by a factor of 21 ± 4 . Here we have used high-resolution *HST*/ACS imaging to build a model of the SPT 0538–50 lens galaxy. By convolving this with the *Spitzer*/IRAC PSF, we then used this model to de-blend the lower-resolution IRAC imaging, separating the emission originating from SPT 0538–50 from that of the foreground lens.

Our main results are as follows.

1. By SED fitting to a range of IR/mm data (included our de-blended IRAC fluxes), we find SPT 0538–50 to have a stellar mass of $(3.3 \pm 1.5) \times 10^{10} M_{\odot}$, and a SFR of $760 \pm 100 M_{\odot} \text{ yr}^{-1}$. This gives an sSFR of 25 Gyr^{-1} , placing SPT 0538–50 a factor of ~ 5 above the main sequence of star-forming galaxies at this epoch, similar to the location (relative to the main sequence) of local merger-driven starbursts. Assuming a simple “closed box” situation, we calculate that SPT 0538–50 will exhaust its molecular gas supply in ~ 20 Myr.
2. We do not detect any far-IR fine structure lines in the spectrum of SPT 0538–50. We attribute this to the presence of a high ionization parameter U .
3. We detect several highly excited molecular emission lines (including H_2O and two high- J lines of ^{12}CO), which, when compared to other similar galaxies, suggest highly excited conditions in the ISM. This could be the signature of highly compact star formation, though boosting via differential lensing may also be contributing to this effect.

The derived physical properties all point toward SPT 0538–50 undergoing a highly compact, merger-driven mode of star formation, comparable to local ULIRGs, and unlike the extended mode of star formation discovered in some other DSFGs at this epoch.

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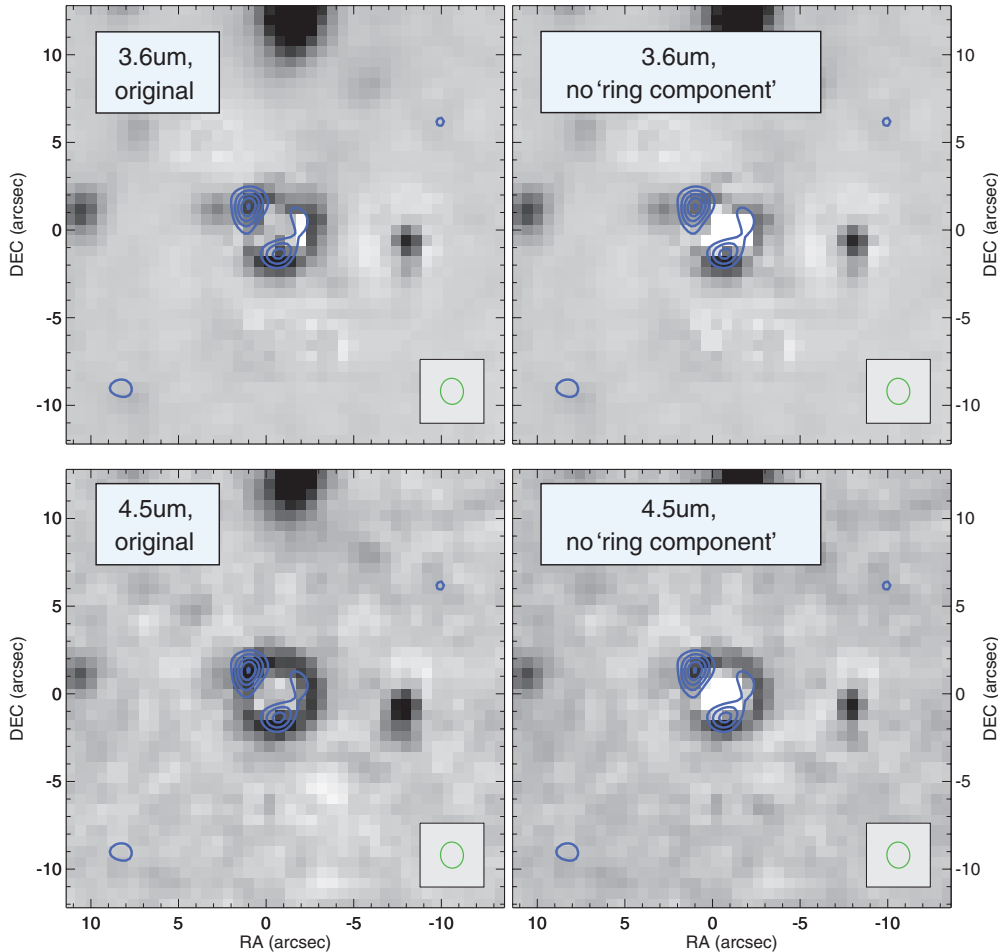


Figure 11. Comparison of IRAC subtraction residuals (at both $3.6 \mu\text{m}$ and $4.5 \mu\text{m}$), obtained with and without the use of a dummy normalizing “ring” component. The image colorscale has been stretched in order to highlight “over-subtraction” issues. In both bands, the subtraction without the use of the ring component is poorer than our original, properly normalized data (presented in Section 3.2.2 and Figure 6) due to over-subtraction in the center of the rings. We include these comparison images to demonstrate that ring-like residuals (which match the ALMA emission) are not artifacts of the normalizing component, but are present even when using more simplistic normalization methods.

(A color version of this figure is available in the online journal.)

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Facilities: *HST/WFC3*, *Spitzer/IRAC*, *ALMA*, *APEX*, *ATCA*, *VLT:Kueyen*, *SMA*, *Herschel*

APPENDIX A

THE LENS GALAXY

Using the remaining IRAC flux (i.e., the flux residual after subtracting the background DSFG), we can also estimate a stellar mass for the foreground elliptical galaxy. Zhu et al. (2010) give the following prescription for calculating the stellar mass of a galaxy from the IRAC $3.6 \mu\text{m}$ luminosity:

$$\log M_* = (-0.79 \pm 0.03) + (1.19 \pm 0.01) \log \nu L_\nu[3.6 \mu\text{m}]. \quad (\text{A1})$$

With the contribution from SPT 0538–50 subtracted, the foreground lens has a $3.6 \mu\text{m}$ flux of $S_{3.6} = 382 \pm 65 \mu\text{Jy}$.

We K-correct this flux to the rest frame $3.6 \mu\text{m}$ using the prescription appropriate for early type galaxies given by Huang et al. (2007), which results in a (rest-frame) luminosity of $\nu L_\nu = (2.6 \pm 0.4) \times 10^{10} L_\odot$. Using Equation (A1), this then gives a stellar mass of $(3.9 \pm 0.8) \times 10^{11} M_\odot$.

Although the foreground lens galaxy of SPT 0538–50 is certainly a massive elliptical in a somewhat crowded environment, it is unlikely that it is the Brightest Cluster Galaxy (BCG) of a small cluster. As can be seen from our *HST/WFC3* imaging (see Figure 6, top right panel), the associated structure extends less than 100 kpc from the lensing galaxy, which is far smaller than the typical extent of even a small cluster. Combined with the relatively modest stellar mass (for a BCG), this suggests that SPT 0538–50 represents a galaxy-lensed system rather than a cluster lens.

APPENDIX B

DECONVOLUTION WITHOUT THE RING-SHAPED NORMALIZING COMPONENT

As we discuss in Section 3.2.2, it is important to note that the “ring”-like structure of SPT 0538–50 seen in the IRAC residuals is not a result of using the normalizing ring component (which is

included only to achieve an optimum subtraction). Even crude normalization techniques (such as simply matching the peak flux of the model to the peak flux of the real image) recover ring-shaped residuals that match the ALMA submillimeter continuum.

These cruder methods do not find an optimum subtraction (due to the DSFG emission being spread across the lens galaxy by the IRAC PSF), but it is important to stress that the ring-like residuals do not result from the use of the ring-shaped normalization component.

Figure 11 shows a comparison of the $3.6\ \mu\text{m}$ and $4.5\ \mu\text{m}$ residuals obtained with and without the use of the dummy normalizing component. While the quality of the subtraction is lower in the latter case (with some over-subtraction present in the center of the ring), we still easily recover the ring-like structure associated with SPT 0538–50, demonstrating that it is real.

REFERENCES

- Abel, N. P., Dudley, C., Fischer, J., Satyapal, S., & van Hoof, P. A. M. 2009, *ApJ*, **701**, 1147
- Abel, N. P., Ferland, G. J., Shaw, G., & van Hoof, P. A. M. 2005, *ApJS*, **161**, 65
- Aravena, M., Murphy, E. J., Aguirre, J. E., et al. 2013, *MNRAS*, **433**, 498
- Banerji, M., Chapman, S. C., Smail, I., et al. 2011, *MNRAS*, **418**, 1071
- Baugh, C. M., Lacey, C. G., Frenk, C. S., et al. 2005, *MNRAS*, **356**, 1191
- Benson, A. J. 2012, *NewA*, **17**, 175
- Bothwell, M. S., Smail, I., Chapman, S. C., et al. 2013, *MNRAS*, **429**, 3047
- Bradford, C. M., Aguirre, J. E., Aikin, R., et al. 2009, *ApJ*, **705**, 112
- Burgarella, D., Buat, V., & Iglesias-Páramo, J. 2005, *MNRAS*, **360**, 1413
- Carilli, C., & Walter, F. 2013, *ARA&A*, **51**, 105
- Carlstrom, J. E., Ade, P. A. R., Aird, K. A., et al. 2011, *PASP*, **123**, 568
- Chabrier, G. 2003, *PASP*, **115**, 763
- Clements, D. L., Rigby, E., Maddox, S., et al. 2010, *A&A*, **518**, L8
- Coppin, K. E. K., Chapman, S. C., Smail, I., et al. 2010, *MNRAS*, **407**, L103
- Coppin, K. E. K., Danielson, A. L. R., Geach, J. E., et al. 2012, *MNRAS*, **427**, 520
- Danielson, A. L. R., Swinbank, A. M., Smail, I., et al. 2011, *MNRAS*, **410**, 1687
- Devillard, N. 2001, *ASPC*, **238**, 525
- Elbaz, D., Dickinson, M., Hwang, H. S., et al. 2011, *A&A*, **533**, A119
- Engel, H., Tacconi, L. J., Davies, R. I., et al. 2010, *ApJ*, **724**, 233
- Fazio, G. G., Hora, J. L., Allen, L. E., et al. 2004, *ApJS*, **154**, 10
- Fixsen, D. J., Bennett, C. L., & Mather, J. C. 1999, *ApJ*, **526**, 207
- Frayser, D. T., Ivison, R. J., Scoville, N. Z., et al. 1998, *ApJL*, **506**, L7
- Fu, H., Cooray, A., Feruglio, C., et al. 2013, *Natur*, **498**, 338
- Genzel, R., Tacconi, L. J., Gracia-Carpio, J., et al. 2010, *MNRAS*, **407**, 2091
- González-Alfonso, E., Fischer, J., Isaak, K., et al. 2010, *A&A*, **518**, L43
- Graciá-Carpio, J., Sturm, E., Hailey-Dunsheath, S., et al. 2011, *ApJL*, **728**, L7
- Greve, T. R., Bertoldi, F., Smail, I., et al. 2005, *MNRAS*, **359**, 1165
- Greve, T. R., Vieira, J. D., Weiß, A., et al. 2012, *ApJ*, **756**, 101
- Griffin, M. J., Abergel, A., Abreu, A., et al. 2010, *A&A*, **518**, L3
- Hainline, L. J., Blain, A. W., Smail, I., et al. 2011, *ApJ*, **740**, 96
- Häussler, B., McIntosh, D. H., Barden, M., et al. 2007, *ApJS*, **172**, 615
- Hayward, C. C., Narayanan, D., Kereš, D., et al. 2013, *MNRAS*, **428**, 2529
- Hezaveh, Y. D., Marrone, D. P., Fassnacht, C. D., et al. 2013, *ApJ*, **767**, 132
- Hezaveh, Y. D., Marrone, D. P., & Holder, G. P. 2012, *ApJ*, **761**, 20
- Holland, W. S., Robson, E. I., Gear, W. K., et al. 1999, *MNRAS*, **303**, 659
- Hopwood, R., Wardlow, J., Cooray, A., et al. 2011, *ApJL*, **728**, L4
- Huang, J.-S., Ashby, M. L. N., Barmby, P., et al. 2007, *ApJ*, **664**, 840
- Ibar, E., Ivison, R. J., Cava, A., et al. 2010, *MNRAS*, **409**, 38
- Ivison, R. J., Swinbank, A. M., Swinyard, B., et al. 2010, *A&A*, **518**, L35
- Kamenetzky, J., Glenn, J., Rangwala, N., et al. 2012, *ApJ*, **753**, 70
- Karim, A., Schinnerer, E., Martínez-Sansigre, A., et al. 2011, *ApJ*, **730**, 61
- Kennicutt, R. C., Jr. 1998, *ARA&A*, **36**, 189
- Kovács, A., Chapman, S. C., Dowell, C. D., et al. 2006, *ApJ*, **650**, 592
- Kroupa, P. 2001, *MNRAS*, **322**, 231
- Lacey, C. G., Baugh, C. M., Frenk, C. S., et al. 2010, *MNRAS*, **405**, 2
- Luhman, M. L., Satyapal, S., Fischer, J., et al. 2003, *ApJ*, **594**, 758
- Lupu, R. E., Scott, K. S., Aguirre, J. E., et al. 2012, *ApJ*, **757**, 135
- Ma, C., Arias, E. F., Eubanks, T. M., et al. 1998, *AJ*, **116**, 516
- Makovoz, D., & Marleau, F. 2005, *PASP*, **117**, 1113
- Meijerink, R., Kristensen, L. E., Weiß, A., et al. 2013, *ApJL*, **762**, L16
- Michałowski, M. J., Dunlop, J. S., Cirasuolo, M., et al. 2012, *A&A*, **541**, A85
- Mocanu, L. M., Crawford, T. M., Vieira, J. D., et al. 2013, arXiv:1306.3470
- Moorwood, A., Cuby, J.-G., Biereichel, P., et al. 1998, *MsAgr*, **94**, 7
- Murphy, E. J., Stierwalt, S., Armus, L., Condon, J. J., & Evans, A. S. 2013, *ApJ*, **768**, 2
- Naylor, B. J., Ade, P. A. R., Bock, J. J., et al. 2003, *Proc. SPIE*, **4855**, 239
- Negrello, M., Hopwood, R., De Zotti, G., et al. 2010, *Sci*, **330**, 800
- Neri, R., Genzel, R., Ivison, R. J., et al. 2003, *ApJL*, **597**, L113
- Noll, S., Burgarella, D., Giovannoli, E., et al. 2009, *A&A*, **507**, 1793
- Omont, A., Neri, R., Cox, P., et al. 2011, *A&A*, **530**, L3
- Omont, A., Yang, C., Cox, P., et al. 2013, *A&A*, **551**, 115
- Panuzzo, P., Rangwala, N., Rykala, A., et al. 2010, *A&A*, **518**, L37
- Papadopoulos, P. P., & Greve, T. R. 2004, *ApJL*, **615**, L29
- Peng, C. Y., Ho, L. C., Impey, C. D., & Rix, H.-W. 2010, *AJ*, **139**, 2097
- Petrov, L., Phillips, C., Bertarini, A., Murphy, T., & Sadler, E. M. 2011, *MNRAS*, **414**, 2528
- Pilbratt, G. L., Riedinger, J. R., Passvogel, T., et al. 2010, *A&A*, **518**, L1
- Rangwala, N., Maloney, P. R., Glenn, J., et al. 2011, *ApJ*, **743**, 94
- Riechers, D. A., Carilli, C. L., Maddalena, R. J., et al. 2011, *ApJL*, **739**, L32
- Rigby, E. E., Maddox, S. J., Dunne, L., et al. 2011, *MNRAS*, **415**, 2336
- Sault, R. J., Teuben, P. J., & Wright, M. C. H. 1995, in *ASP Conf. Ser.* 77, *Astronomical Data Analysis Software and Systems IV*, ed. R. A. Shaw, H. E. Payne, & J. J. E. Hayes (San Francisco, CA: ASP), 433
- Schuster, M. T., Marengo, M., & Patten, B. M. 2006, *Proc. SPIE*, **6270**, 65
- Serjeant, S. 2012, *MNRAS*, **424**, 2429
- Siringo, G., Kreysa, E., Kovács, A., et al. 2009, *A&A*, **497**, 945
- Smail, I., Ivison, R. J., & Blain, A. W. 1997, *ApJL*, **490**, L5
- Solomon, P. M., & Vanden Bout, P. A. 2005, *ARA&A*, **43**, 677
- Spinoglio, L., Dasyra, K. M., Franceschini, A., et al. 2012, *ApJ*, **745**, 171
- Swinbank, A. M., Papadopoulos, P. P., Cox, P., et al. 2011, *ApJ*, **742**, 11
- Valtchanov, I., Virdee, J., Ivison, R. J., et al. 2011, *MNRAS*, **415**, 3473
- Vernet, J., Dekker, H., D'Odorico, S., et al. 2011, *A&A*, **536**, A105
- Vieira, J. D., Crawford, T. M., Switzer, E. R., et al. 2010, *ApJ*, **719**, 763
- Vieira, J. D., Marrone, D. P., Chapman, S. C., et al. 2013, *Natur*, **495**, 344
- Walter, F., Weiß, A., Downes, D., Decarli, R., & Henkel, C. 2011, *ApJ*, **730**, 18
- Wardlow, J. L., Cooray, A., De Bernardis, F., et al. 2013, *ApJ*, **762**, 59
- Werner, M. W., Roellig, T. L., Low, F. J., et al. 2004, *ApJS*, **154**, 1
- Yang, C., Gao, Y., Omont, A., et al. 2013, *ApJL*, **771**, L24
- Yun, M. S., Aretxaga, I., Ashby, M. L. N., et al. 2008, *MNRAS*, **389**, 333
- Zhu, Y.-N., Wu, H., Li, H.-N., & Cao, C. 2010, *RAA*, **10**, 329